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Errata: Solar and Heliospheric Processes from Solar Wind Composition Measurements

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ERRATA

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Solar and heliospheric processes from solar wind composition measurements

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The following errors occurred in the editorial process and were not due to the authors.

Figure 5 on p. 222 should appear as follows:

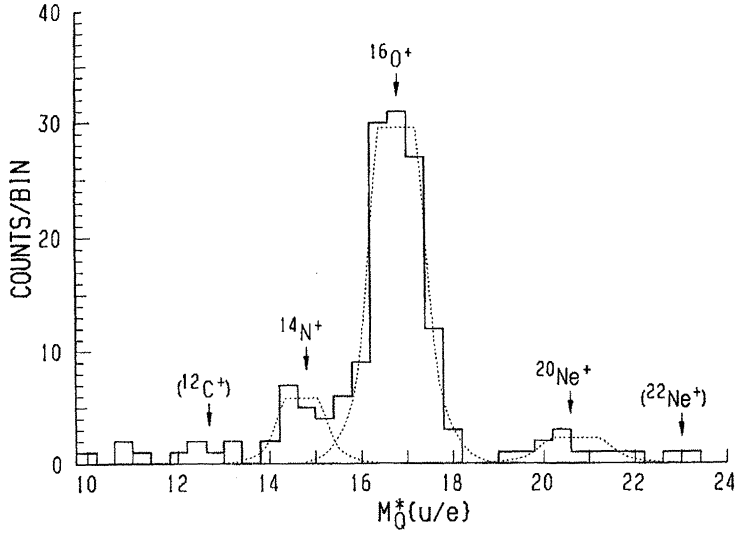

Figure 5. Mass/charge spectrum of heavy ions measured at heliocentric distances from 4.56–5.40 AU (139 days). Both double and triple coincidences are included. MQ^* is the M/Q scale valid for the multiply charged solar wind ions. The exact M/Q ratios for the singly charged interstellar ions, differing from the MQ^* scale, are marked by arrows. Since only ions with $V/V_{SW} > 1.3$ are included the background is very low. N^+, O^+ and Ne^+ were unambiguously identified (Geiss *et al.* 1994*a*).

Figure 6 on p. 224 should appear as follows:

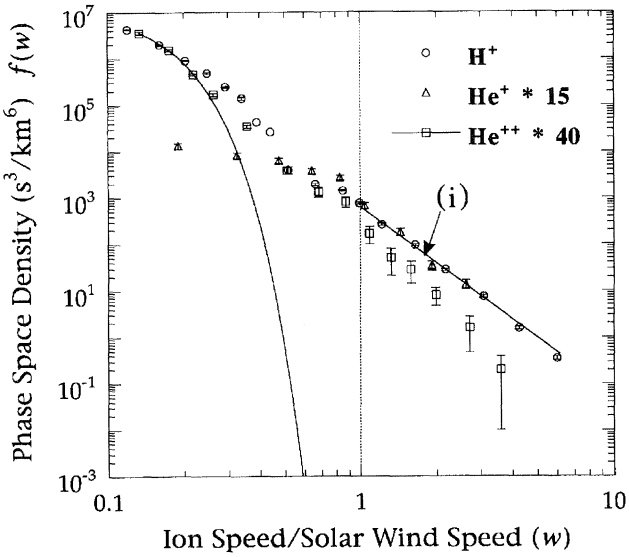

Figure 6. Velocity distributions $f(w)$ of H^+, He^+ and He^{++} as a function of w, the ion speed in the solar wind frame of reference divided by the solar wind bulk speed, observed during a one-day period (Oct 19, 1991) at the leading edge of a corotating interaction region (CIR). The spectra of He^+ and He^{++} have been normalized to that of H^+ at $w = 1$ and $w = 0$ respectively. Clear evidence for the preferential acceleration of the non-thermal pick-up He^+ is the presence of the power law tail, $f(w) = f_0 w^{-4}$, at speeds beyond $w = 1$ where interstellar He^+ is about ten times more abundant than the suprathermal tail of solar wind He^{++}, although the latter has a much higher local density.

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